

First detections of 610 MHz radio emission from hot magnetic stars

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ABSTRACT

We have carried out a study of radio emission from a small sample of magnetic O- and B-type stars using the Giant Metrewave Radio Telescope, with the goal of investigating their magnetospheres at low frequencies. These are the lowest frequency radio measurements ever obtained of hot magnetic stars. The observations were taken at random rotational phases in the 1390 and the 610 MHz bands. Out of the 8 stars, we detect five B-type stars in both the 1390 and the 610 MHz bands. The O-type stars were observed only in the 1390 MHz band, and no detections were obtained. We explain this result as a consequence of free-free absorption by the free-flowing stellar wind exterior to the closed magnetosphere. We also study the variability of individual stars. One star - HD 133880 - exhibits remarkably strong and rapid variability of its low frequency flux density. We discuss the possibility of this emission being coherent emission as reported for CU Vir by Trigilio et al. (2000).

Key words: radiation mechanisms:non-thermal — stars: massive — stars: individual: HD 133880, etc. — stars: magnetic field — radio continuum: stars

1 INTRODUCTION

The last decade has witnessed the identification and elaboration of an important population of hot O- and B-type stars hosting strong, organized magnetic fields (e.g. Donati et al. 2002; Oksala et al. 2010; Grunhut et al. 2012; Wade et al. 2012). Observational and theoretical studies of the interaction of their intense radiation-driven winds with their magnetic fields (e.g. ud-Doula & Owocki 2002; Townsend et al. 2005; Sundqvist et al. 2012) have revealed that the stellar wind properties of magnetic hot stars are modified in important ways as compared to the winds of non-magnetic stars, introducing significant changes across the entire electromagnetic spectrum.

The presence of an organized magnetic field at the surface of a hot star leads to channelling and confinement of its outflowing stellar wind, creating a magnetosphere, which can radiate in various wavebands (Andre et al. 1988; Linsky et al. 1992).

In regions close to the star, the magnetic pressure dominates and causes the wind to follow dipolar field lines. At greater distances from the stellar surface, the wind kinetic energy density

exceeds the magnetic pressure due to the stronger radial decline of magnetic field energy density than of wind kinetic energy density. The radius at which the wind kinetic energy density becomes equal to the magnetic pressure is defined as the Alfvén radius (e.g. ud-Doula & Owocki 2002). The region below the Alfvén radius, i.e. interior to the largest closed magnetic loop, is defined as the inner magnetosphere (Trigilio et al. 2004), which is the site of generation of X-ray and H α emission (e.g. Babel & Montmerle 1997; Gagné et al. 2005; Howarth et al. 2007; Sundqvist et al. 2012). Beyond the Alfvén radius, the wind opens the magnetic field lines, generating a current sheet in the magnetic equatorial plane. This region is defined as the middle magnetosphere. At greater distances from the star, the outer magnetospheric region is characterized by a radial magnetic topology following wind streamlines.

Radio emission from non-magnetic hot stars is expected to be thermal free-free emission from the ionized wind in the circumstellar environment (Panagia & Felli 1975; Wright & Barlow 1975). However, in the presence of magnetic fields electrons can be accelerated to relativistic energies, either by reconnection near the current sheet in the middle magnetosphere (Usov & Melrose 1992) or generated in strong, large-scale shocks in the inner magnetosphere (Babel & Montmerle 1997; ud-Doula et al. 2006, 2008). These en-

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ergetic electrons can give rise to gyrosynchrotron radio emission (Drake et al. 1987; Linsky et al. 1992; Triglilio et al. 2004).

Various surveys of cooler magnetic A- and B-type stars have been carried out at 5–8 GHz radio frequencies, resulting in $\approx 25\%$ detections (e.g. Abbott et al. 1985; Drake et al. 1987; Linsky et al. 1992). Linsky et al. (1992) interpreted the radio emission in these stars as gyrosynchrotron in nature, arising from regions located at $10\text{--}20 R_*$, with higher frequency emission originating from plasma located closer to the star, and the lower frequency emission from material further away.

Leone & Umana (1993) monitored the 5 GHz radio emission of two magnetic B-type stars, HD 37017 and σ Ori E, and observed variability in accordance with their respective rotational periods. The coincidence of radio maxima with the extrema of the longitudinal magnetic field in these stars led those authors to suggest that radio emitting regions are located above the magnetic poles. Radio variability was also measured from the magnetic B8p star HD 133880 (HR Lup). The flux density and circular polarization were observed to vary according to the stellar rotation period (Lim et al. 1996; Bailey et al. 2012), and was also interpreted as gyrosynchrotron emission (Bailey et al. 2012).

Leone et al. (2004) argued that as per the standard radio emission gyrosynchrotron models, emission at a given frequency should be emitted in a well-localised torus above the magnetic poles. They examined the radio spectral energy distributions (SEDs) of five magnetic B-type stars in the range from 1 GHz to 87.7-GHz. They concluded that the observed slopes of the radio spectra and the absence of millimetre emission are not generally compatible with this model, and suggested a cut-off frequency of the radio emission.

Triglilio et al. (2000) detected rapid, intense, narrow-band and highly polarized radio bursts from the late B-type star CU Vir, which they hypothesised to be Electron Cyclotron Maser Emission (ECME) at 1.4 GHz. Such emission is expected to occur principally at low frequencies (< 1.5 GHz, Triglilio et al. 2000; Leto et al. 2012). However, to date, few such low frequency observations have been obtained, and generally they correspond to null results (e.g. George & Stevens 2012).

Clearly, a variety of diverse investigations of the radio emission of magnetic A- and B-type stars have been carried out, leading to an array of disconnected results. However, no similar studies have yet taken place for hotter magnetic O- and B-type stars. In an attempt to extend and homogenize the study of the physics of radio magnetospheres of hot stars, we have initiated a systematic survey of the radio emission properties of the known magnetic B- and O-type stars. Our survey includes observations at both high frequency (currently being obtained with the VLA) and at low frequency (currently being obtained with the GMRT).

In this first paper we describe a pilot study of 8 established magnetic O- and B-type stars, observed with the GMRT in the 1390 MHz (20 cm) and 610 MHz (50 cm) bands, to characterize the radio emission properties of their magnetospheres at very low frequencies, study their radio SEDs to understand magnetospheric properties and probe mass loss from the star in the case of dominant free-free absorption.

In §2 we describe the observations and data analysis. In §3 we carry out a detailed analysis and look for variability of the detected stars within the observing period, and in §3.6 we discuss the stars for which no radio emission was detected. In §4 we discuss our results and the potential of low frequency radio observations for providing new constraints on the physics of hot star magnetospheres.

2 OBSERVATIONS

The GMRT observations of 8 magnetic O- and B-type stars were taken between 2013 October 24 to 2014 January 24 under cycle 25 in the 1390 (L-band) and 610 MHz bands. Table 1 gives details of the observed stars. Out of these, 5 B-type stars were observed in both the 1390 and 610 MHz bands, while the 3 O-type stars were observed only in the 1390 MHz band. Each observation was of approximately 3 hours in duration and a total of 26–29 good antennae were used. For all observations, visibilities were recorded for two circular polarizations (RR and LL) with a bandwidth 32 MHz, divided into 256 frequency channels, and 16 s integration time. Calibrator sources were used to remove the effect of variation of the instrumental and other non-astronomical factors in the measurements. 3C48, 3C286 and 3C147 were used as flux calibrators in various observations. Flux calibrators were observed for 10–15 minutes, either in the beginning or towards the end of each observation. Flux calibrators were also used for bandpass calibration. Phase calibrators were chosen from the VLA calibrator manual such that they are located within 15 degrees of the target star. They were observed more frequently. This is important not only for the tracking of instrumental phase and gain drifts and atmospheric and ionospheric gain and phase variations but also for monitoring the quality and sensitivity of the data and for spotting occasional gain and phase jumps. In the 1390 MHz band, the phase calibrator observations were made approximately every 25–30 minutes for 5 minutes duration, whereas in the 610 MHz band, the 6-minute phase calibrator scans were obtained every 35–40 minutes.

A FLAGging and CALibration (FLAGCAL) software pipeline developed for automatic flagging and calibration of the GMRT data (Prasad & Chengalur 2012) was used to flag and calibrate the GMRT data. In 70% of the cases, the results were satisfactory and only minor additional flagging was needed after the FLAGCAL. However, in some cases, the data needed more significant flagging and recalibration. In such cases the flagging was done manually and a flag file was created. The FLAGCAL was then run again using the manual flag file. The calibrated data were inspected to determine the presence of any spurious signature. The above step was repeated until satisfactory calibration was achieved. The flagged and calibrated visibility data were used to make continuum images using the standard tasks in the Astronomical Image Processing System (AIPS). To avoid bandwidth smearing in the continuum image, the total bandwidth was divided into 6 sub-bands in the 1390 MHz observations, and in 10 sub-bands in the 610 MHz observations. Because of the large field of view (FoV) of the GMRT, the three-dimensional (3D) imaging feature in the AIPS task ‘IMAGR’ was used in which the entire FoV is divided into multiple subfields (facets) and each of which was imaged separately. For the 1390 MHz image, the total FoV was divided in 19 subfields, and in the 610 MHz band the FoV was divided into 37 subfields. The presence of a large number of bright sources in the FoV of various stars allowed us to carry out self-calibration to improve the complex gains. This reduces the errors from temporal variations in the system gain, and spatial and temporal variations in the ionospheric properties. After 3 rounds of phase self-calibration, the clean component model was subtracted from the UV data to identify the residual bad data. Some more flagging was performed and 2 more rounds of phase self-calibration were run. A final round of amplitude and phase self-calibration was also performed in each dataset. Since the phase variations occur on time-scales of a few minutes, the time interval for the phase self-calibration was chosen to be 1 minute, and 5 minutes for the amplitude and phase self-calibration.

Table 1. Details (spectral type, effective temperature, luminosity, radius, mass, rotational period, polar strength of the magnetic dipole at the stellar surface, Kepler co-rotation radius, Alfvén radius, mass loss rate and wind terminal velocity) of magnetic O and B stars observed with the GMRT. R_{ff} is the radius of the free-free absorption photosphere at 1390 MHz, discussed in Sect. 3.6.

Star	Spectral Type	T_{eff} (kK)	$\log L_{\star}$ (L_{\odot})	R_{\star} (R_{\odot})	M_{\star} (M_{\odot})	Period (d)	B_p (kG)	R_K (R_{\star})	R_A (R_{\star})	$\log(\dot{M})$ ($M_{\odot} \text{ yr}^{-1}$)	v_{∞} (km s $^{-1}$)	R_{ff} (R_{\star})	R_{ff}/R_A
HD215441	B8-9p	15 ± 2	1.8	2.6	3.5	9.49	34	-	-	-	-	-	-
HD37479	B2 Vp	23 ± 2^1	3.6^1	3.9^1	5.0^1	1.19^2	9.6^3	2.1^3	31^3	-9.774^6	1794^6	6.5	0.2
HD37017	B2 Vp	21 ± 2^4	3.4^4	3.9^4	7.2^4	0.90^4	$> 6.0^5$	1.9^6	$> 18^6$	-9.072^6	1102^6	27.6	< 1.5
HD36485	B3 Vp	20 ± 2^7	3.5^7	4.5^7	7.1^7	1.48^7	10^7	2.4^6	24^6	-8.969^6	1012^6	30.4	1.3
HD133880	B8IVp	13 ± 1^8	2.0^8	2.0^8	3.2^8	0.88^8	19^8	2.2^8	60^8	-11.0^8	750^8	4.6	0.1
HD37022	O7 Vp	39 ± 1^9	5.3^9	9.9^9	45^9	15.42^{10}	1.1^{11}	9.4^6	2.4^6	-6.399^6	3225^6	236.3	98.5
HD57682	O9 V	34 ± 1^{12}	4.8^{12}	7.0^{12}	17^{12}	63.57^{13}	1.7^{13}	24^6	3.7^6	-7.079^6	2395^6	153.7	41.5
NGC1624-2	O6.5f?cp	35 ± 2^{14}	5.1^{14}	9.7^{14}	34^{14}	158.0^{14}	$> 20^{14}$	41^{14}	$> 11^{14}$	-6.786^6	2890^6	151.2	< 13.7

¹Hunger et al. 1989, ²Townsend et al. 2010, ³Oksala et al. 2012, ⁴Bolton et al. 1998, ⁵Bohlender et al. 1987, ⁶Petit et al. 2013, ⁷Leone et al. 2010, ⁸Bailey et al. 2012, ⁹Simon-Diaz et al. 2006, ¹⁰Stahl et al. 2008, ¹¹Wade et al. 2006, ¹²Grunhut et al. 2009, ¹³Grunhut et al. 2012, ¹⁴Wade et al. 2012

Table 2. GMRT observations of magnetic B and O type stars.

Star	Obs Date 1.4 GHz	Mean HJD	Mean Phase	Flux density F_{20} (mJy)	rms μ Jy	Obs Date 0.6 GHz	Mean HJD	Mean Phase	Flux density F_{50} (mJy)	rms μ Jy
HD 215441	01-Nov-13	2456598.262± 0.050	0.93	1.49 ± 0.10	46	24-Oct-13	2456590.168± 0.044	0.08	0.98 ± 0.10	56
HD 37479	31-Oct-13	2456597.433± 0.170	0.09	2.00 ± 0.11	57	24-Oct-13	2456590.402± 0.044	0.18	1.25 ± 0.28	51
HD 37017	31-Oct-13	2456597.422± 0.131	0.94	1.59 ± 0.12	53	24-Oct-13	2456590.430± 0.040	0.18	0.59 ± 0.32	153
HD 36485	27-Oct-13	2456593.378± 0.050	0.94	1.30 ± 0.10	54	26-Oct-13	2456591.494± 0.076	0.66	0.59 ± 0.17	88
HD 133880	24-Jan-14	2456681.616± 0.055	0.84	14.45 ± 0.11	49	14-Jan-14	2456671.590± 0.048	0.42	2.04 ± 0.15	66
HD 37022	31-Oct-13	2456597.476	0.4	< 4.95	1650
HD 57682	01-Nov-13	2456598.370	0.13	< 0.16	52
NGC 1624-2	02-Nov-13	2456598.487	1.0	< 0.21	71

The map resolutions for the 1390 MHz images were around 2''–3'', and around 5''–6'' for the 610 MHz data. In Table 2, we report the values of the final flux density. For non-detections, we quote 3σ upper limits.

3 ANALYSIS AND RESULTS

With our GMRT observations, we have detected 5 out of 8 stars. The detected stars, HD 215441, HD 37479, HD 37017, HD 36485 and HD 133880, are B-type stars and have been detected in both the 1390 and 610 MHz bands. The O-type stars, HD 37022, HD 57682 and NGC 1624-2, were observed only in the 1390 MHz band, and none of them are detected.

Of the B stars, 4 stars (HD 37479, HD 37017, HD 36485 and HD 133880) are sufficiently rapid rotators that rotation contributes significantly to the support of their magnetospheres (so-called ‘centrifugal magnetospheres’; Petit et al. 2013). On the other hand, the O-type stars, as well as HD 215441, rotate slowly, and their magnetospheres receive no significant rotational support (so-called ‘dynamical magnetospheres’; Petit et al. 2013).

The observations are summarized in Table 2. In this section, we give details of the detected B-type stars, and explore potential causes of non-detections in O-type stars.

3.1 HD 215441

HD 215441 (Babcock’s star) is a cool magnetic B8-9p star. It exhibits no optical, UV or X-ray evidence of a magnetosphere, but was detected in the 5 GHz and 1.4 GHz bands (flux density ~ 1 mJy) by Linsky et al. (1992). Despite its low temperature, this

star is noteworthy due to its extraordinarily strong magnetic field (Babcock 1960).

Magnetic O- and B-type stars show variability of many observational quantities. This variability is understood as rotational modulation in the context of the oblique rotator model (ORM) (Stibbs 1950; Wade et al. 2011), in which the dipolar surface magnetic field is tilted with respect to the rotational axis. Rotation of the star results in strict periodic variability according to the stellar rotational period.

To phase our observations of HD 215441, we used the ephemeris of North & Adelman (1995):

$$\text{JD} = 2448733.714 + (9.487574 \pm 0.000030) \cdot E. \quad (1)$$

Radio emission is clearly detected in both the 1390 and 610 MHz bands at phases 0.93 and 0.08, respectively, with a flux density somewhat larger than that reported at higher frequencies by Linsky et al. (1992, 1.2–1.3 mJy). In both bands, our observations cover less than 1% of the rotational period; hence we are unable to evaluate any variability. The flux is somewhat ($\sim 15\%$) lower in the 610 MHz band. Whether this is due to the slope of the radio SED, or due to flux variability between phases 0.93 and 0.08, is not currently known.

Linsky et al. (1992) reported a flux of 1.07 ± 0.09 mJy at 1.4 GHz on July 2, 1987. According to Eq. 1, this corresponds to phase 0.08, i.e. the same phase as our 610 MHz observation (0.98 ± 0.1 mJy). Comparing the two measurements, this implies an essentially flat spectral index of ($\alpha = \log(F_{1390}/F_{610})/\log(1390/610) = +0.02 \pm 0.03$) at this phase. This result appears inconsistent with the $\alpha = -0.26$ index derived by Leone et al. (2004) based on higher frequency observations. This suggests that the SED flattens at low frequency.

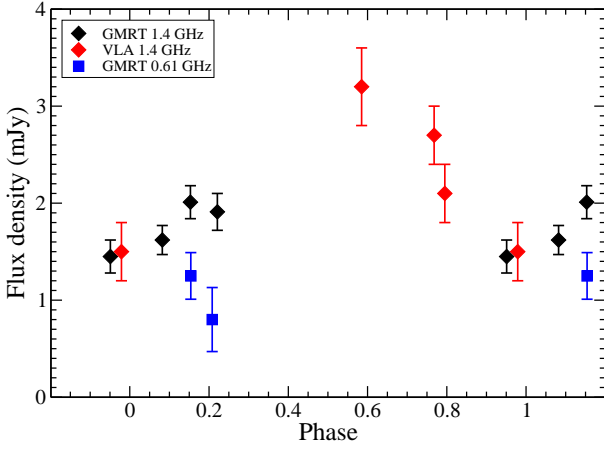


Figure 1. GMRT and VLA 1.4 GHz measurements (black and red diamonds, respectively) and GMRT 610 MHz measurements (blue squares) of HD 37479 (σ Ori E). The VLA measurements are taken from the archive.

3.2 HD 37479

HD 37479 (σ Ori E) is the prototypical B2 Vp star hosting a centrifugal magnetosphere. It exhibits strong and variable $H\alpha$ emission that has been studied in some detail (e.g. Townsend et al. 2005; Oksala et al. 2012). It was detected by Linsky et al. (1992) at 15, 5 and 1.4 GHz frequency bands using the VLA, and observed at several phases. The flux density was observed to vary by a factor of about 2, from 2.4–3.1 mJy at 15 GHz, 2.8–3.9 mJy at 5 GHz, and 1.5–3.2 mJy at 1.4 GHz bands. They also detected variable circular polarization at 15 GHz.

This star was also observed by Leone & Umana (1993), who detected 5 GHz emission varying with the rotational period, with maxima coinciding with the magnetic extrema.

We obtained multiple GMRT observations in both bands. To phase the data, we use the ephemeris of Townsend et al. (2010), which takes into account the (slow) rotational braking of the star:

$$JD = 2442778.8290 + 1.1908229 \cdot E + 1.44 \times 10^{-9} \cdot E^2. \quad (2)$$

Our observations of HD 37479 were obtained at phases 0.95, 0.08, 0.13 and 0.22 in the 1390 MHz band, and phases 0.15 and 0.21 in the 610 MHz band. We have supplemented our observations with the four 1.4 GHz measurements obtained by Linsky et al. (1992). The phased flux variation is shown in Fig. 1.

The combined 1390 MHz data suggest a roughly sinusoidal variation peaking at phase 0.5, with a peak flux of ~ 3.2 mJy. The 610 MHz measurements are both weaker than the minimum 1390 MHz measurement, but their phase coverage is too sparse to infer meaningful information about the phase variation in this band.

Leone & Umana (1993) proposed that the 5 GHz flux from this star exhibited a double-wave phase variation. This is not suggested to be the case at 1.4 GHz band. Nevertheless, additional data will be required to draw more robust conclusions.

Two of the 1390/610 MHz GMRT measurements were obtained at similar phases. From these we infer an instantaneous ($\phi = 0.15 - 0.2$) index of the SED at these frequencies of $\alpha = 0.11 \pm 0.04$. This is consistent with $\alpha = 0.12$ derived by Leone et al. (2004) based on higher frequency observations.

3.3 HD 37017

HD 37017 is an early-type magnetic B2 Vp star that also exhibits $H\alpha$ emission due to the presence of a dense centrifugal magnetosphere (Leone et al. 2010; Petit et al. 2013). This target is in fact a double-lined spectroscopic binary (SB2) system with an 18.65d orbital period containing the magnetic B star in addition to a late B dwarf (Leone et al. 2010).

Leone & Umana (1993) detected variable radio emission from HD 37017. Like σ Ori E, the radio emission was found to vary with the rotational period, and magnetic extrema coincided with radio maxima. It was detected by Linsky et al. (1992) at 15, 5 and 1.4 GHz bands using the VLA, and observed at several phases. The flux density was observed to vary by a factor of about 2, from 0.9–2.1 mJy at 15 GHz, 1.4–2.6 mJy at 5 GHz, and 1.5–2.4 mJy at 1.4 GHz bands. Leone et al. (2004) studied HD 37017 in mm bands and examined the properties of the radio spectrum in the 1.4–87.7 GHz range. They found the flux density to increase up to a frequency of 22.5 GHz and then to decrease in the mm range. This indicates a possible cut-off frequency.

We obtained GMRT observations in the 1390 and 610 MHz bands. We phased the data using the ephemeris of Bolton et al. (1998):

$$JD = 2446010.3750 + (0.9011836 \pm 0.000006) \cdot E \quad (3)$$

Our 1390 MHz observations were obtained at phases 0.07, 0.80 and 0.98, while our 610 MHz observations were obtained at phases 0.15 and 0.21. We supplemented these measurements with the 1420 MHz observations of Linsky et al. (1992). The phased data are illustrated in Fig. 2.

Leone & Umana (1993) described the 5 GHz flux variation as sinusoidal, with a maximum at phase 0.0 and a minimum at phase 0.5. The variation shown in Fig. 2, although weak and coarsely sampled, appears opposite to this description. This is entirely a consequence of the different ephemeris used by those authors and by us. The phasing relative to the longitudinal field variation is the same.

Again, our two 610 MHz observations exhibit lower flux than the lowest of the 7 measurements at 1390 MHz. One of the VLA 1.4 GHz observations was obtained at a phase similar to the GMRT 610 MHz data. From these we infer an instantaneous ($\phi = 0.15 - 0.2$) index of the SED at these frequencies of $\alpha = 0.23 \pm 0.1$. This is consistent with $\alpha = 0.15$ derived by Leone et al. (2004) based on higher frequency observations.

3.4 HD 36485

HD 36485 (δ Ori C) is another hot B3 Vp star hosting a centrifugal magnetosphere. The $H\alpha$ emission is variable on a timescale of a few hours. Notwithstanding its sharp spectral lines, the star is a rapid rotator with a period of about 1.5d, implying that it is viewed close to the rotational pole (Bohlender et al. 1991; Leone et al. 2010).

δ Ori C was detected with the VLA at 5 GHz with a mean flux density of 0.95 mJy and showed a negative spectral index ($\alpha < -0.7$; $f_\nu \propto \nu^\alpha$) between 5 and 15 GHz bands according to Drake et al. (1987), who concluded a non-thermal origin of the radio emission. Their 3 measurements at this wavelength did not show any significant variability, with a measurement uncertainty of ± 0.1 mJy.

We use the ephemeris of Leone et al. (2010):

$$JD = 2448298.86 + (1.47775 \pm 0.00004) \cdot E. \quad (4)$$

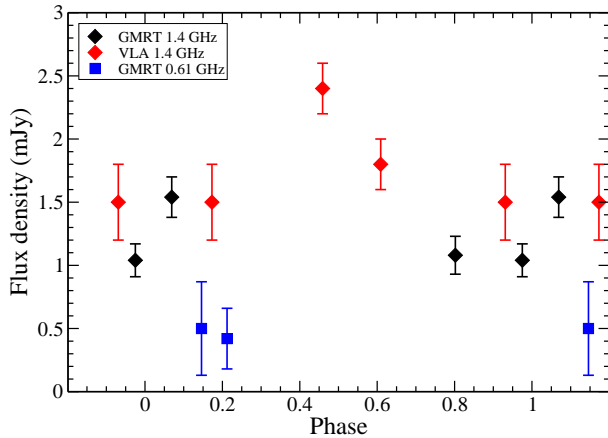


Figure 2. GMRT and VLA 1.4 MHz measurements (black and red diamonds, respectively) and GMRT 610 MHz measurements (blue squares) of HD 37017. The VLA measurements are taken from the archive.

In our GMRT observations, we detected this star in both the 1390 and 610 MHz bands. In our 3 hour observation at 1390 MHz covering phases 0.91–0.96, the flux density varied between 0.42 ± 0.24 to 0.6 ± 0.14 mJy. Although these measurements are formally consistent with a constant flux density, the flux did monotonically increase during the observation, suggesting a real change. In the 610 MHz band, the flux density varied significantly, from 1.46 ± 0.16 to 0.69 ± 0.13 mJy between phases 0.62 to 0.71.

3.5 HD 133880

HD 133880 (HR Lup) is a cooler, rapidly rotating ($P = 0.88$ d; Bailey et al. 2012) Bp star hosting one of the strongest known stellar magnetic fields. However, unlike most Ap/Bp stars, HD 133880 has a magnetic field topology that appears to be predominantly quadrupolar as opposed to dipolar (Landstreet 1990; Bailey et al. 2012). HD 133880 was previously observed with the Australia Telescope Compact Array (ATCA) at 5 and 8.5 GHz frequency bands simultaneously on 1995 February 12, 14 and 16, for around 10 hours each day. Lim et al. (1996) had demonstrated that radio flux and circular polarization of HD 133880 at both frequencies vary significantly and coherently according to the rotational period. They reported that the emission shows broad peaks near the phases of the longitudinal field extrema. The reanalysis of the data by Bailey et al. (2012) yielded the same result. They found flux variation to be complex, characterized by strong, broad maxima at phases 0.0 and 0.5 (i.e. the extrema of the longitudinal field), and sharper, somewhat weaker secondary extrema at quadrature phases (i.e. 0.25 and 0.75).

George & Stevens (2012) observed this star with the GMRT in dual mode in the 610 and 240 MHz bands on December 5 and 7th, 2009. They reported non-detections in all the observations. However, we extracted the data from the GMRT archive and reanalysed it, and find detections on both days in the 610 MHz band. Furthermore, we also find variability in each observation in this band. To re-confirm our detections of their data, we have analysed the data both in AIPS, as well as in Common Astronomy Software Applications (CASA), separating out LL and RR Stokes, and we find consistent results. However, in the 240 MHz band we obtained no detection in either observation, consistent with George & Stevens (2012). The $3 - \sigma$ upper limits in this band at the two epochs were 1.5 and 3.0 mJy, respectively. We also located three data sets at

1420 MHz in the VLA archive, obtained between 15–17 February 1995. We reanalyse these data and tabulate them in Table 3, along with the flux density in various time intervals in both our new observations, as well as the reanalysed observations.

To phase the data, we used the ephemeris of Bailey et al. (2012):

$$JD = 2445472.0 + (0.877476 \pm 0.000009) \cdot E. \quad (5)$$

The GMRT observations were made in the 1390 and 610 MHz bands. In Fig. 3, we plot the phased data in both bands. The flux variation of HD 133880 at both 1390 and 610 MHz is extraordinary. Unlike the other stars observed in this program (Table 4) that exhibit a flux variation of about a factor of 2, the 610 MHz flux of HD 133880 varies by a factor of 16 in our observations (and more than a factor of 12 at 1390 MHz). Even with the sparse phase coverage of our observations, it is clear that the maxima of both the 610 and 1390 MHz data occur at phases ~ 0.25 and 0.75 , coincident with the phases of the minor maxima at 5 and 8 GHz described by Bailey et al. (2012).

Considering the important differences in the phase variations of the GMRT data and the ATCA data, we speculate that HD 133880 may exhibit maser emission in a manner similar to CU Vir (Trigilio et al. 2000). In CU Vir, while no coherent emission was found at 5 GHz, the 1.4 GHz emission was identified as ECME with a basal flux of 2–3 mJy and then very large increments in the flux density at around phases 0.35–0.45 and 0.75–0.85. While, we do not yet have complete phase coverage for HD 133880, the observed features are highly reminiscent of those of CU Vir (see Fig. 3).

3.6 Free-free absorption in O star winds

Schnerr et al. (2007) carried out 5 and 1.4 GHz observations of 5 (non-magnetic) O stars with the Westerbork Synthesis Radio Telescope (WSRT). They detected 3 stars, for the first time at 1.4 GHz band, with one star (ζ Per) inferred to show non-thermal radio emission while two other (α Cam and λ Cep) were inferred to show thermal radio emission. In all cases, the observed flux was lower than their predicted thermal flux, based on the mass loss rates inferred from the H α line profiles. They explained this discrepancy as due to a higher influence of wind clumping on the formation of the H α line.

In our program we also observed 3 magnetic O-type stars at 1390 MHz. None of these targets were detected (see Table 2). For one of these stars (HD 37022) our upper limit is quite poor because it resides in a region with extended radio emission in Orion.

The high mass loss rates of O stars can make the wind optically thick due to free-free absorption out to a very large “radio photosphere”. In such a case radio observations are unlikely to directly see the magnetospheric emission. On the other hand, they may probe the global mass loss, which is predicted to be quite heavily reduced for O stars with dynamical magnetospheres, due to quenching of the mass loss by the field and the resultant plasma in-fall (ud-Doula & Owocki 2002; ud-Doula et al. 2008).

We can estimate the radius of the free-free radio photosphere, under the assumption of a spherically-symmetric, non-magnetic wind, using Eq. (4) of Torres (2011):

$$\tau_{ff} = 5 \times 10^3 \dot{M}^2 v_\infty^{-2} f^{-1} \nu^{-2} T^{-3/2} R_{ff}^{-3} \quad (6)$$

where τ_{ff} is the free-free optical depth, \dot{M} is the mass-loss rate in units of $10^{-8} M_\odot/\text{yr}$, v_∞ is the wind terminal velocity in units of 10^8 cm/s, f is the clump volume filling factor, ν is the frequency of

Table 3. Variability of HD 133880 .

Freq GHz	Telescope	Date	Mean HJD	Phase	Flux density mJy	rms μ Jy
1.388	GMRT	24-Jan-14	2456681.572 \pm 0.012	0.791	25.90 \pm 0.14	108
		24-Jan-14	2456681.604 \pm 0.012	0.827	26.14 \pm 0.12	123
		24-Jan-14	2456681.636 \pm 0.012	0.863	3.41 \pm 0.12	117
		24-Jan-14	2456681.663 \pm 0.008	0.895	2.43 \pm 0.14	177
1.425	VLA	15-Feb-95	2449764.040 \pm 0.007	0.348	2.07 \pm 0.40	396
		15-Feb-95	2449764.098 \pm 0.006	0.414	2.87 \pm 0.33	331
		16-Feb-95	2449765.013 \pm 0.006	0.456	3.39 \pm 0.24	239
		17-Feb-95	2449766.010 \pm 0.006	0.593	5.86 \pm 0.27	267
0.608	GMRT	14-Jan-14	2456671.552 \pm 0.010	0.371	1.91 \pm 0.23	76
		14-Jan-14	2456671.579 \pm 0.010	0.402	1.97 \pm 0.24	68
		14-Jan-14	2456671.607 \pm 0.010	0.434	2.21 \pm 0.24	65
		14-Jan-14	2456671.631 \pm 0.007	0.461	1.91 \pm 0.14	75
0.606	GMRT	05-Dec-09	2455170.712 \pm 0.026	0.966	2.37 \pm 0.31	311
		05-Dec-09	2455170.776 \pm 0.027	0.038	2.16 \pm 0.33	328
0.607	GMRT	07-Dec-09	2455172.663 \pm 0.027	0.189	1.26 \pm 0.32	218
		07-Dec-09	2455172.726 \pm 0.026	0.261	20.08 \pm 0.40	287
		07-Dec-09	2455172.787 \pm 0.024	0.330	3.52 \pm 0.37	347

Table 4. Variability of other stars in our sample.

Star	Tel.	Freq GHz	Date	Mean HJD	Phase	Flux density mJy	rms μ Jy	Ref.
HD215441	GMRT	1.388	01 Nov 13	2456598.223 \pm 0.010	0.927	1.43 \pm 0.18	76	This paper
			01 Nov 13	2456598.255 \pm 0.010	0.931	1.15 \pm 0.15	70	
			01 Nov 13	2456598.283 \pm 0.010	0.934	1.45 \pm 0.18	74	
			01 Nov 13	2456598.306 \pm 0.005	0.936	1.20 \pm 0.19	92	
HD215441	GMRT	0.608	02 Jul 87	2446979.250	0.078	1.07 \pm 0.09	...	This paper
			24 Oct 13	2456590.136 \pm 0.012	0.075	1.19 \pm 0.13	71	
			24 Oct 13	2456590.168 \pm 0.012	0.078	1.18 \pm 0.13	69	
			24 Oct 13	2456590.201 \pm 0.012	0.082	0.70 \pm 0.22	108	
HD37479	GMRT	1.388	31 Oct 13	2456597.273 \pm 0.010	0.951	1.45 \pm 0.17	87	This paper
			31 Oct 13	2456597.430 \pm 0.010	0.082	1.62 \pm 0.15	84	
			01 Nov 13	2456597.514 \pm 0.010	0.153	2.01 \pm 0.17	86	
			01 Nov 13	2456597.595 \pm 0.008	0.221	1.91 \pm 0.19	106	
	VLA	1.4	11 Mar 85	2446136.467	0.585	3.2 \pm 0.4	...	Linsky et al. (1992)
			12 Mar 85	2446136.685	0.768	2.7 \pm 0.3	...	
			16 Mar 85	2446141.480	0.795	2.1 \pm 0.3	...	
			17 Mar 85	2446141.699	0.979	1.5 \pm 0.3	...	
HD37479	GMRT	0.608	24 Oct 13	2456590.370 \pm 0.012	0.154	1.25 \pm 0.24	134	This paper
			24 Oct 13	2456590.434 \pm 0.012	0.208	0.80 \pm 0.33	158	
HD37017	GMRT	1.388	31 Oct 13	2456597.301 \pm 0.010	0.802	1.08 \pm 0.15	74	This paper
			31 Oct 13	2456597.458 \pm 0.010	0.975	1.04 \pm 0.13	66	
			01 Nov 13	2456597.542 \pm 0.010	0.069	1.54 \pm 0.16	78	
			01 Nov 13	2456597.542 \pm 0.010	0.069	1.54 \pm 0.16	78	
	VLA	1.4	11 Mar 85	2456136.488	0.459	2.4 \pm 0.2	...	Linsky et al. (1992)
			12 Mar 85	2456136.623	0.609	1.8 \pm 0.2	...	
HD37017	GMRT	0.608	16 Mar 85	2456141.419	0.931	1.5 \pm 0.3	...	This paper
			17 Mar 85	2456141.637	0.173	1.5 \pm 0.3	...	
			24 Oct 13	2456590.402 \pm 0.012	0.146	0.50 \pm 0.37	74	
			24 Oct 13	2456590.462 \pm 0.008	0.212	0.42 \pm 0.24	66	
			27 Oct 13	2456593.338 \pm 0.010	0.910	0.43 \pm 0.17	90	
			27 Oct 13	2456593.365 \pm 0.010	0.929	0.44 \pm 0.33	85	
HD36485	GMRT	1.388	27 Oct 13	2456593.393 \pm 0.010	0.947	0.60 \pm 0.14	92	This paper
			27 Oct 13	2456593.419 \pm 0.009	0.965	0.52 \pm 0.17	78	
			27 Oct 13	2456593.419 \pm 0.009	0.965	0.52 \pm 0.17	78	
			27 Oct 13	2456593.419 \pm 0.009	0.965	0.52 \pm 0.17	78	
	VLA	1.4	07 Jul 86	2446619.810	0.221	< 1.2	400	Linsky et al. (1992)
			25 Oct 13	2456591.429 \pm 0.012	0.618	1.46 \pm 0.16	108	
			25 Oct 13	2456591.461 \pm 0.012	0.640	1.26 \pm 0.18	218	
			26 Oct 13	2456591.526 \pm 0.012	0.684	0.96 \pm 0.19	87	
HD36485	GMRT	0.608	26 Oct 13	2456591.558 \pm 0.012	0.706	0.69 \pm 0.13	97	This paper
			26 Oct 13	2456591.558 \pm 0.012	0.706	0.69 \pm 0.13	97	

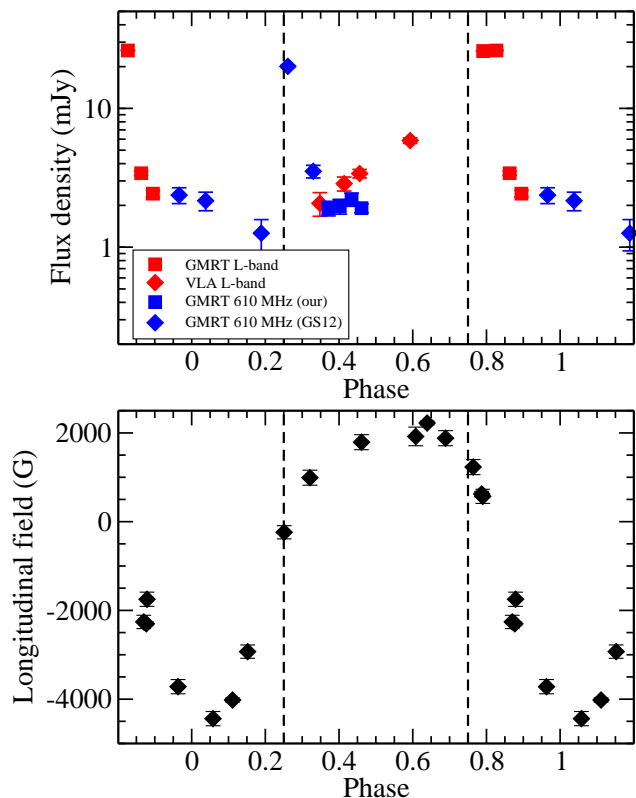


Figure 3. *Upper frame* - Phase versus flux density plot of HD 133880 at 1.4 GHz (red color) and 610 MHz (blue color). Here red squares denote GMRT 1390 MHz measurements, red diamonds denote VLA 1.4 GHz measurements, blue squares denote GMRT 610 MHz observations made by us in 2014 January, whereas the blue diamonds are the 2009 December data of George & Stevens (2012) reanalyzed by us. Note that unlike in Figs. 1 and 2, the vertical scale in this figure is logarithmic. *Lower frame* - Longitudinal magnetic field variation of HD 133880, from Bailey et al. (2012). The vertical lines in both plots indicate phases 0.25 and 0.75.

observation in GHz, T is the wind temperature in units of 10^5 K, which we here for simplicity approximate with the effective temperature of the star, and R_{ff} is the distance from the star, in units of 3×10^{12} cm.

We adopt a volume filling factor $f = 1$, i.e. a smooth, unclumped radio emitting region, since it has been demonstrated (Puls et al. 2006) that $f = 1$ yields the best agreement with observations of thermal radio emission from hot stars when the Vink, de Koter, & Lamers (2000) scaling law for mass-loss rates is assumed. Assuming $\tau_{\text{ff}} = 1$, we have computed the position of the radio photosphere (D , in units of R_*) at 1390 MHz, using the wind parameters (computed using the Vink scaling) and physical parameters reported by Petit et al. (2013), except for HD 133880, for which we use values reported by Bailey et al. (2012). The results are summarized in Table 1 and illustrated in Fig. 4.

If $R_{\text{ff}} < R_A$ we expect the magnetospheric emission to escape the wind and to be detectable (in principle). However, if $R_{\text{ff}} > R_A$, the magnetospheric emission will be hidden within the (thermal) radio photosphere of the wind.

In our sample, B stars show ratios of R_{ff}/R_A which are clustered around values 0.1–1.5, whereas all the O-type stars have this ratio clustered between 10–100. This is illustrated in Fig. 4. This convincingly demonstrates that the free-free absorption of the magnetospheric radio emission of the B stars likely is negligible, and

that the principal cause of the null results for O stars is free-free scattering in the free streaming wind.

Current models of the magnetically confined winds of slowly rotating O stars imply that they host “dynamical magnetospheres”. In this scenario, wind streamlines inside R_A in the magnetic equatorial region are constrained to follow closed field loops. After the wind stream reach the respective loop top and shock, the cooling plasma returns to the stellar surface on a relatively short (dynamical) timescale. Simulations (ud-Doula & Owocki 2002) predict that as a consequence of this confinement and fall-back, the global mass loss of O stars with dynamical magnetospheres should be significantly reduced, by up to a factor of 10–20 compared to similar non-magnetic stars.

In the context of free-free absorption, this reduction in mass loss has important implications for the location of the radio photosphere. The expected mass-loss rate can be estimated using equation 23 of ud-Doula et al. (2008), leading to a reduction of \dot{M} by a factor of more than 3 for HD 37022, 5 for HD 57682, and 15 for NGC 1624-2. If we reduce the mass loss rate in Eq. (6) by these factors, for HD 37022 and HD 57682 the ratio R_{ff}/R_A is reduced to roughly 40, i.e. still well outside the Alfvén radius. However, the ratio for NGC 1624-2 is reduced to ~ 2 . At shorter wavelengths (3–6 cm), this value becomes < 1 . Hence multiwavelength radio observations spanning the radio spectrum have the potential to provide novel constraints on predictions of mass-loss quenching of strongly magnetic O stars like NGC 1624-2.

4 SUMMARY AND DISCUSSION

We have studied 8 O- and B-type stars with the GMRT in the 1390 and 610 MHz bands. Our 610 MHz detections of the 5 B type stars are the first detections of this class of objects at such low frequencies. This indicates that free-free absorption of their winds is not able to shadow their low frequency emission. We also reanalysed archival observations of HD 133880 obtained in December 2009, for which George & Stevens (2012) claimed non-detections. Surprisingly, we do detect the radio emission in both their observations allowing us to diagnose variability.

For HD 37017 and HD 37479, we confirm weak variability of the 1390 MHz flux density, as reported by Leone & Umana (1993) in 5 GHz band. The variability appears to be in phase with the variation of the longitudinal magnetic field. We do not find evidence of a double-wave 1390 MHz variation of HD 37479 (as reported by Leone & Umana 1993, at 5 GHz), although better phase coverage is necessary to robustly confirm this conclusion.

For the B star HD 36485, our observations suggest a relatively rapid change in the flux density in both bands in narrow phase ranges. This result needs to be confirmed by further observations.

The cooler B star HD 133880 shows remarkable variability in both the 1390 and 610 MHz bands. Even with our limited phase coverage, the initial results suggest that the fluxes in both bands peak at phases coinciding with the minor peaks obtained from ATCA data (Bailey et al. 2012). We propose that this behaviour may be a consequence of ECME, first observed from the B8p star CU Vir by Triglio et al. (2000). The phenomenon, only observable at low radio frequencies, strongly motivates immediate monitoring of this target.

There were no detections obtained for any of the three magnetic O stars in the 1390 MHz band. Using the scaling law of Torres (2011), we are able to conclude that this null result is in general agreement with the expected free-free optical depth of dense

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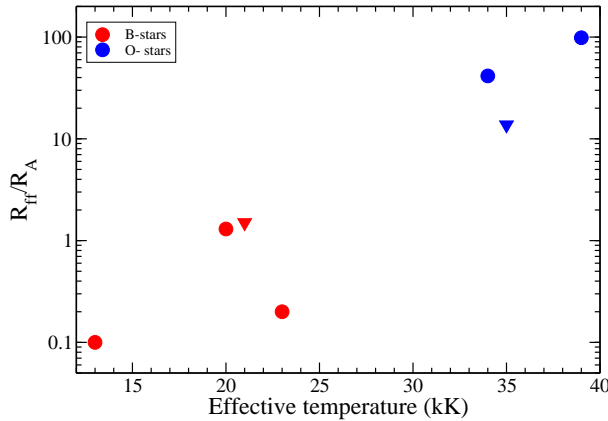


Figure 4. Ratio of R_{ff} to R_A versus effective temperature for the observed sample (excluding HD 215441, for which R_A is unavailable). Here red circles and blue circles are for B- and O- type stars. The inverted triangle indicate upper limits as listed in Table 1.

winds of O-type stars. We also conclude that radio observations at higher frequencies will be capable of testing predictions of mass loss quenching of strongly magnetic O stars like NGC 1624-2.

Leone et al. (2004) commented that the radio spectra of HD 142301 and HD 215441 (the two stars in their sample with the strongest magnetic fields) exhibited a sharp drop in flux density between 15 and 22.5 GHz. They concluded that this suggests there is a cut-off frequency in this range. On the other hand, the magnetic chemically peculiar stars with relatively weak magnetic fields showed flux densities that certainly decrease from the cm to the mm range. In our sample, we were able to measure spectral indices between 1390 and 610 MHz for 3 stars in common with the sample of Leone et al. (2004). For two stars (HD 37017 and HD 37479), the indices were similar to those obtained at higher frequencies, suggesting a smooth decrement of the flux density. On the other hand, for HD 215441 (the nondegenerate star with the strongest known magnetic field), we obtained a spectral index that differed strongly from that derived by Leone et al. (2004). This implies an equally significant change in the spectral index at low frequencies, and deserves further investigation.

The results presented in this paper represent a first feasibility study of the detection of low frequency emission from magnetic O- and B-type stars. Having demonstrated the general properties of a range of objects, we are now carrying out a larger survey to observe all known magnetic OB stars in radio bands from low to high frequencies. We are also performing monitoring of individual magnetic stars, in order to better understand the rotational variation of their fluxes and spectral energy distributions.

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